The case for dark matter: wealth of indirect evidence but \*conclusive\* direct detection remains elusive

\* Galactic rotation curves, gravitational lensing, CMB, BAO, large-scale structure observations and simulations, Bullet Cluster multi-spectral overlay study, nucleosynthesis

Direct searches for dark matter using noble liquids, in particular xenon in recent years, have obtained the best sensitivities in the field for moderate to high-mass dark-matter Weakly Interacting Massive Particles

XENON10, XENON100, ZEPLIN, LUX, PandaX, XENON1T, LZ

\* High Z and density (good self-shielding), scalable (liquid), two-channel energy reconstruction, many advantages

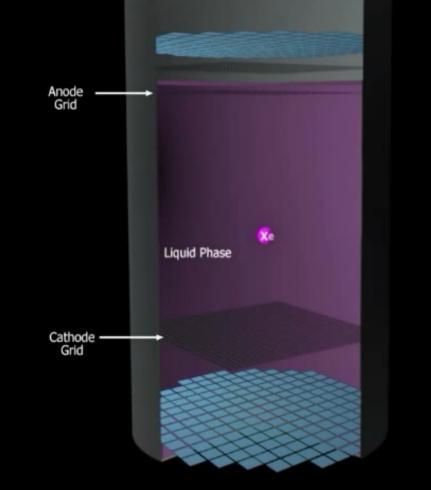
Along with the development of this technology, there has been a continued effort to better understand the detailed scintillation and ionization responses

### \*The Motivation

- \*The dual-phase xenon detector, an example of a time-projection chamber (TPC), with many PMTs
- \*S1 (primary) and S2 (secondary) scintillation, the latter from charge
  - \* Their ratio discriminates between nuclear and electron recoil (NR & ER)
  - \* The sum gives you energy
- \* Fiducialization with multiple-scattering rejection powerful: WIMPs will exclusively lead to single-scatter recoils



WIMP Signals in a Dual-Phase Xenon Detector



Animation credit: The LUX collaboration, but created prior to commencement of detector operations, so the values of parameters are approximate and representative Penning quenching (NR: worst at high energy)

 $E_0 \xrightarrow{N_{ex}} N_{ph} \xrightarrow{} S1 \text{ (geometric light collection times quartz VUV transmission times PE conversion probability,...)} N_i \longrightarrow N_{e-} \xrightarrow{} Drift, diffuse, die \longrightarrow Extraction \xrightarrow{} Gas photons \xrightarrow{} S2$ 

Elastic scattering a.k.a. atomic motion a.k.a. heat (NR: worst at low energy)

\*Chain reaction set off by just 1 NR or ER leads to many NRs and ERs, with 3 main processes occurring

\*Billiard-ball scattering, electron excitation, and ionization

- \*Working theory of the physics leading to scintillation (light collection) and escaping ionization electrons (charge collection)
- \*Focusing heavily upon liquid xenon for today's talk
  - \*But, gaseous xenon, argon, and other noble elements and phases work within same general framework



- \*Concepts are incorporated into NEST (Noble Element Simulation Technique) which is a Monte Carlo tool
- \*Recent improvements in our understanding found in B. Lenardo et al., arXiv:1412:4417 and on the NEST web sites
  - \*<u>http://www.albany.edu/physics/NEST.shtml</u>

\*<u>http://nest.physics.ucdavis.edu/site/</u>

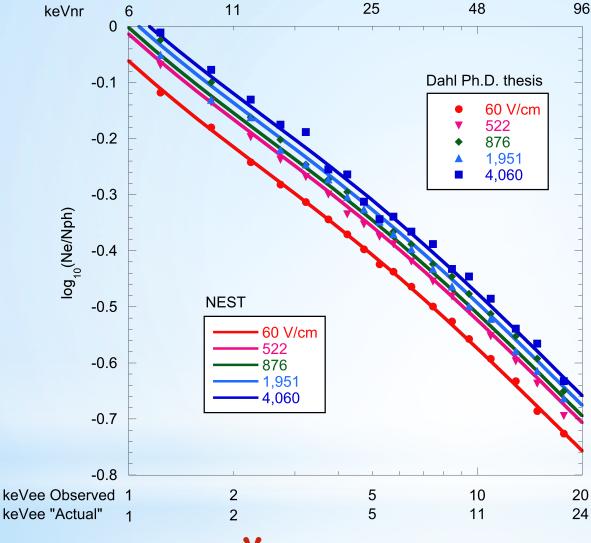
- \*The older NEST papers
  - \*J. Mock et al., JINST 9 (2014) T04002. <u>arXiv:1310.1117</u>
  - \*M. Szydagis et al., JINST 8 (2013) C10003. <u>arXiv:1307.6601</u>
  - \*M. Szydagis et al., JINST 6 (2011) P10002. <u>arXiv:1106.1613</u>
- \*Examples of both postdictive and predictive power of approach presented here today
- \*Free parameters used at every step have physical meaning





\*W = 13.7 + -0.2 eV (higher than xenon excitation or ionization potentials because of cases where electrons fail to jump levels) \* Determined empirically from energy scale of ERs, combining the reconstructed photon and electron counts, just like in Germanium \*Sources: C.E. Dahl thesis, Princeton 2009, and many other works \*If  $N_{ex}/N_i = 0.15$  for ER (best-fit NEW model, 0.06 outdated) then \*Observe the unification and simplification: Traditional  $W_i = E_0 / N_i$ =  $[(N_{ex} + N_i) * W]/N_i = (N_{ex}/N_i + 1) * W = 1.15 * 13.7 eV ~ 15.8 eV$ \*Compare to Takahashi 1975 result of 15.6 +/- 0.3 eV; others similar \*This is not forced: pieces fit naturally with a new understanding \*S1 and S2,  $N_{ph}$  and  $N_{e-}$ ,  $N_{ex}$  and  $N_i$ , these are all (see Conti 2003) anti-correlated. For ER this means ~fixed sum. NR is tricky (L)

E. Conti et al., Phys. Rev. B68 (2003) 054201 T. Takahashi et al., Phys. Rev. A12 (1975) 1771

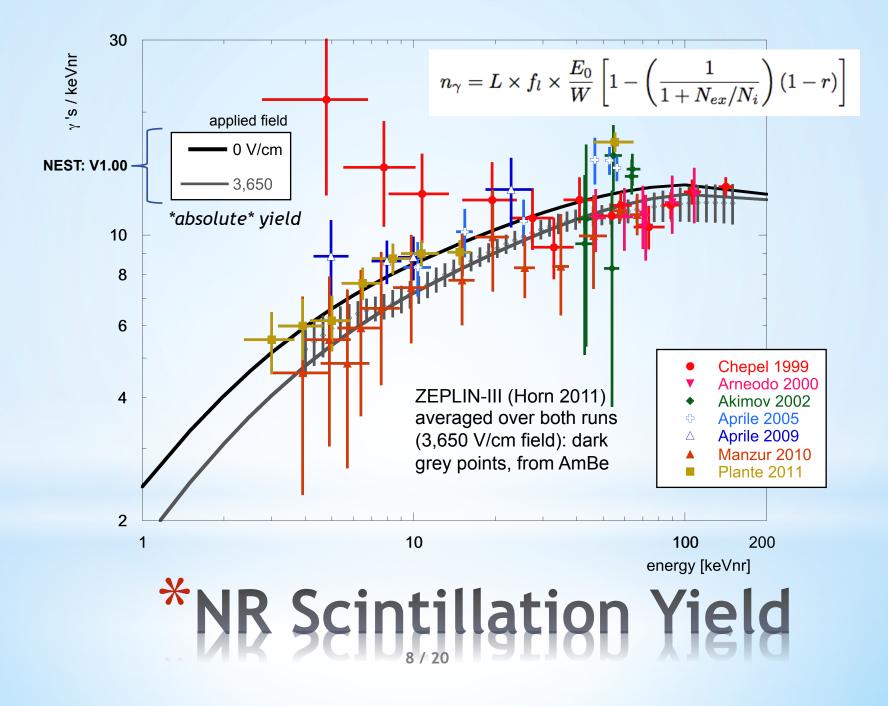


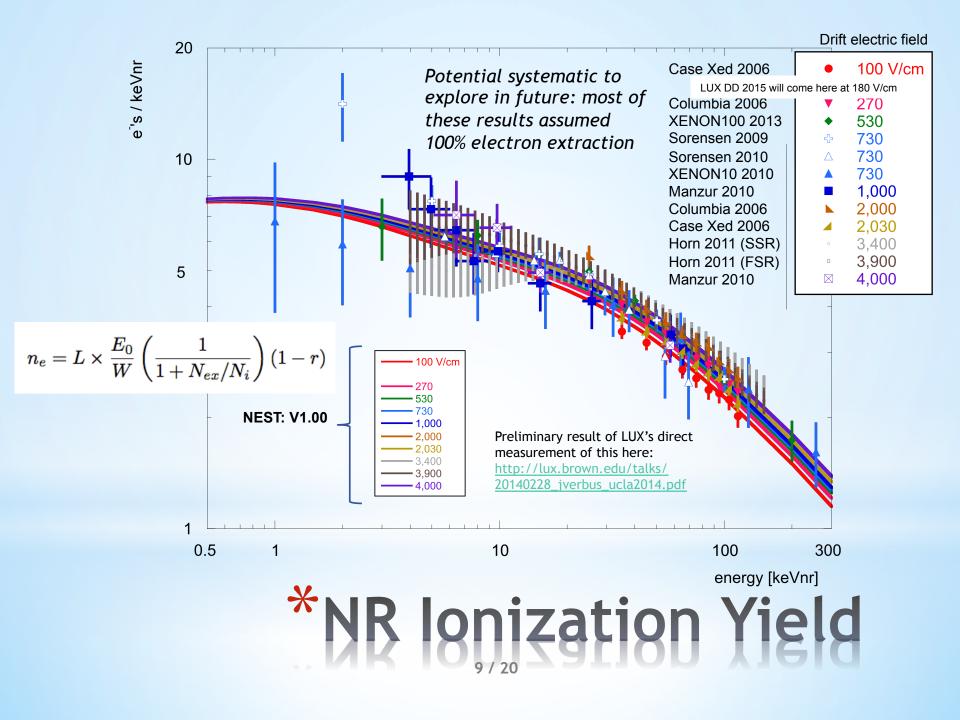
$$E_{nr} = \mathcal{L}^{-1} \cdot (n_e + n_\gamma) \cdot W.$$

The keVnr (or keVr) energy scale is an estimate of the actual energy of the recoil for NR. keVee is same but without a L(indhard) factor

The keVee (or keVer) energy scale (or "electron equivalent") is an estimate of the actual energy of the recoil for ER. For NR, it is the ER for which the average total number of quanta is the same (traditionally used only the S1, but this created a field-dependent energy)

## \*NR Charge/Light Ratio





Parameter	Best Fit	68% conf.	
α	1.240	+0.079	
		-0.073	
ζ	0.0472 +0.0088		
		-0.0073	
β	239 +28		
		-8.8	
$\gamma$	0.01385	+0.00058	
		-0.00073	
δ	0.0620	+0.0056	
		-0.0064	
k	0.1394	+0.0032	
		-0.0026	
η	3.3	+5.3	
		-0.06	
λ	1.14	+0.45	
		-0.09	
С	0.00555	+0.00025	
		-0.00025	

# By the Numbers $N_{ex}/N_i = \alpha F^{-\zeta} \left(1 - e^{-\beta \epsilon}\right)$

\* Reduction in free parameters over older NEST - Less splining and more physical motivation

 $\varsigma = \gamma F^{-o}$  \* Combined fit of light, charge simultaneously

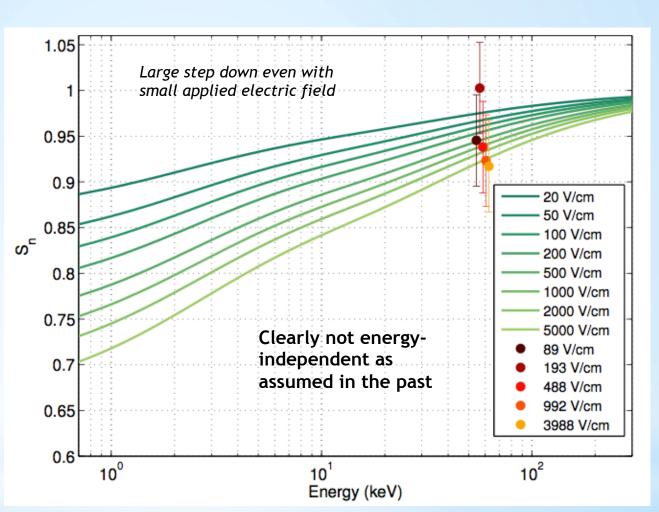
\* Global fit over as much data as possible - Moving far beyond Dahl data

quenching of photon yield (high dE/dx)

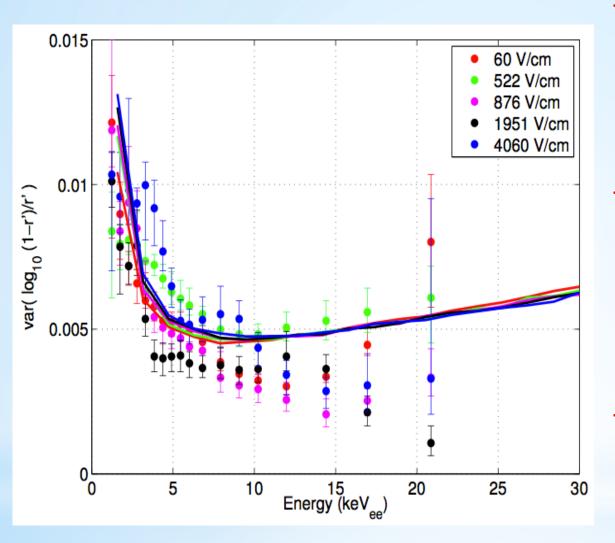
\* Over-conservativeness (low yields) removed

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- \*S<sub>n</sub> is scintillation light yield relative to zero E-field
- \*Only useful direct data at 56.5 keVnr (Aprile 2005)
- \*Continuity of zero and non-zero field models achieved with effective minimum field
- \*Light yield goes down as charge up (more escape)



# \* Relative Yield x. Field

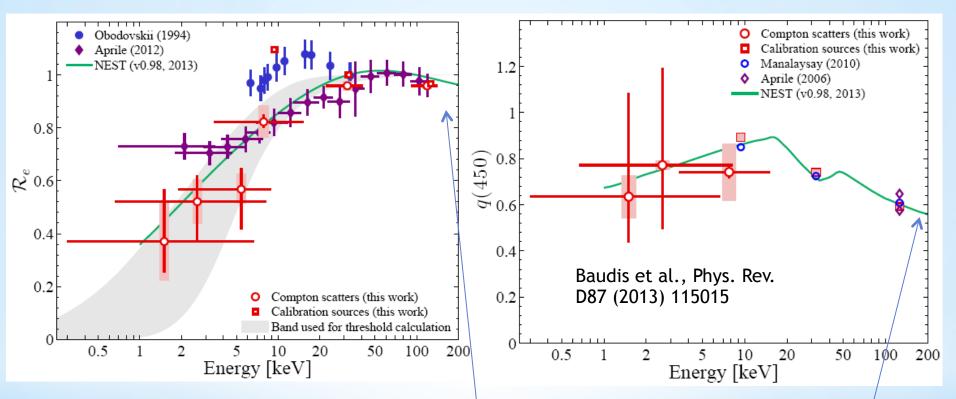


- \*Energy resolution in xenon long known not to be that expected from binomial recombination
- \*Solution is to utilize a a special variablewidth Poisson function
  - \* Compromise between unphysical Gaussian and a slow binomial
- \*Recombination Fanolike factor proportional to N<sub>i</sub> (A. Dobi thesis)

# \* NR Band Width

#### Zero applied electric field

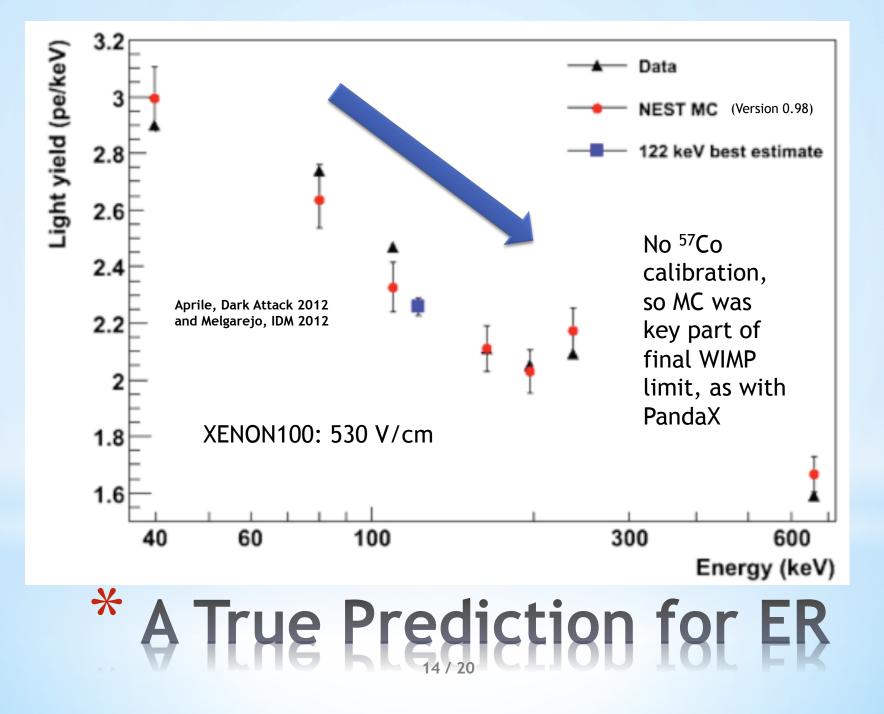
#### Non-zero field (450 V/cm)

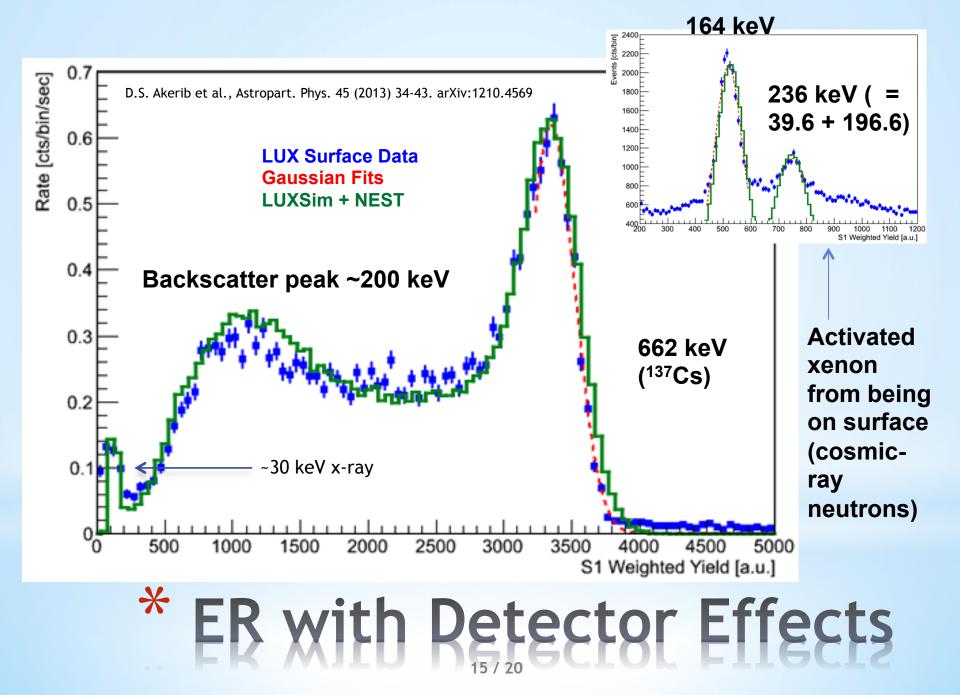


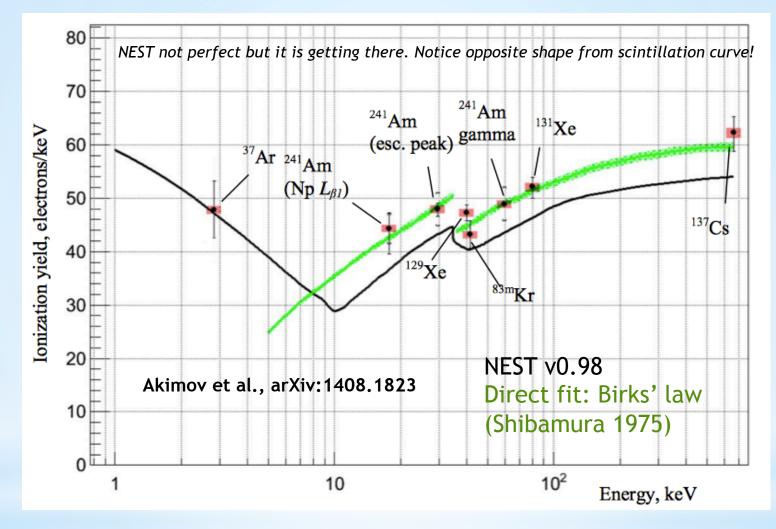
At high energies, recombination and thus yields behave differently

For ER, fit to electron data only and gammas/x-rays must follow: simpler approach

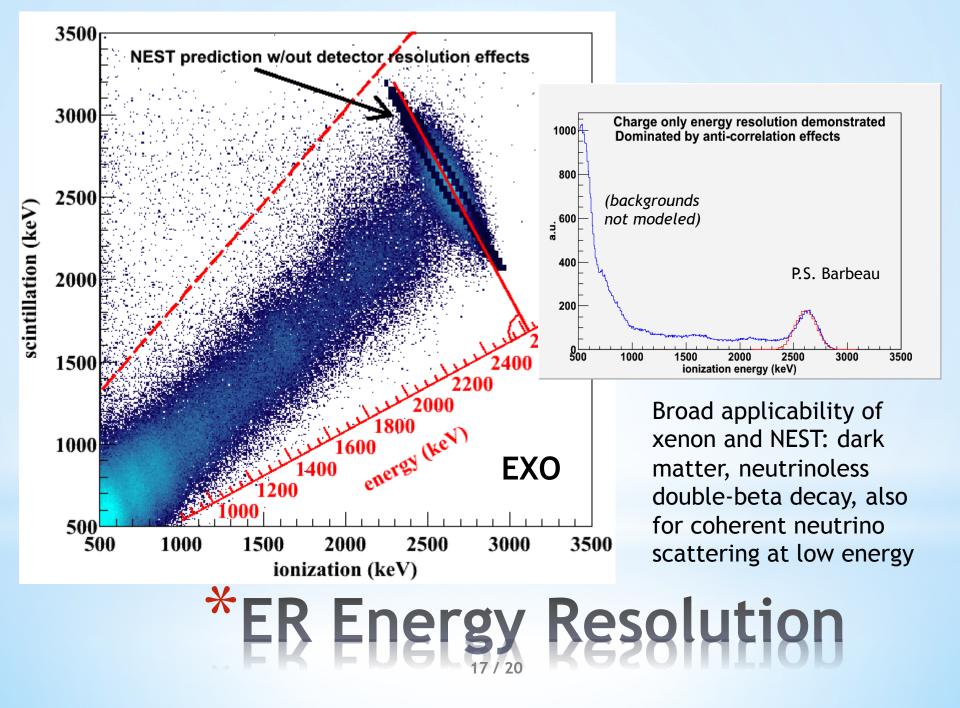
### \* ER Scintillation Yield

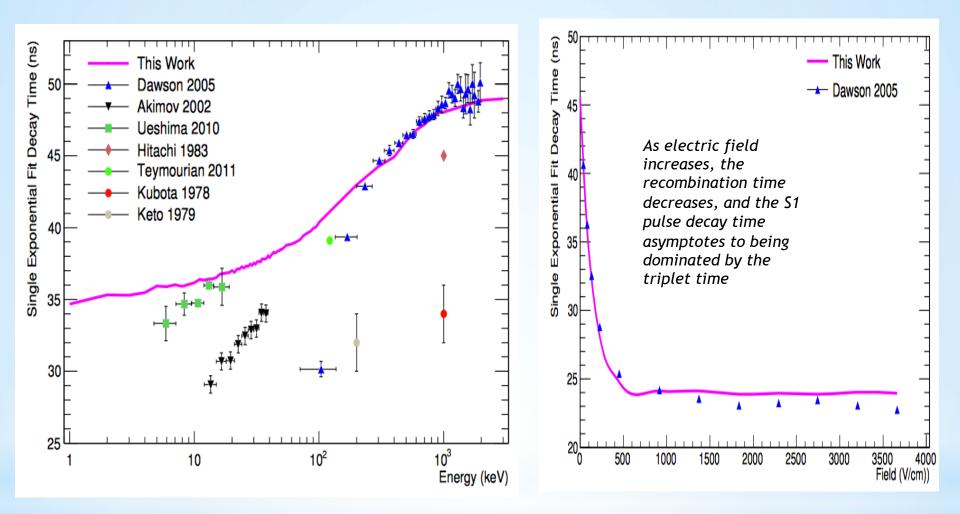




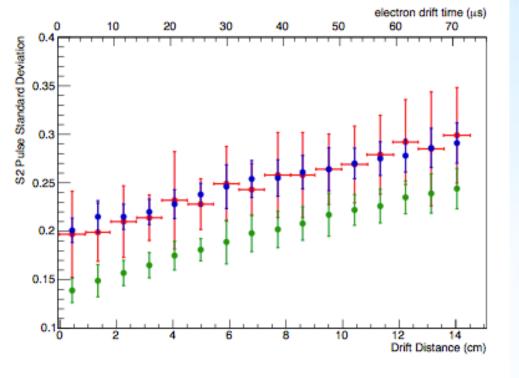


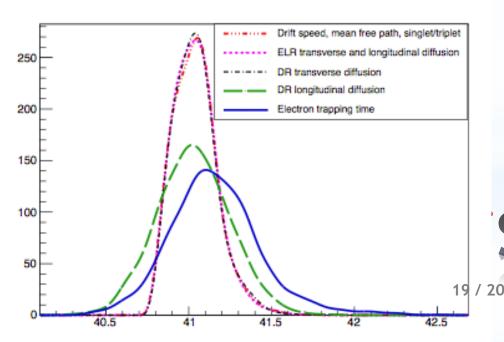
## \*ER Ionization Yield





Singlet lifetime	$3.1\pm0.7~\mathrm{ns}$			
Triplet lifetime	$24 \pm 1$ ns		Dulas	Change
Singlet/Triplet - ER from direct excitation ( $\gamma$ induced)	$0.17\pm0.05$		PIIISe	Shape
Singlet/Triplet - ER from recombination ( $\gamma$ induced)	$0.8\pm0.2$		I MISE	SIGRE
Singlet/Triplet - ER from both processes ( $\alpha$ induced)	$2.3 \pm 0.51$		DIICO	<b>Nana</b>
Singlet/Triplet - NR (neutron induced)	$7.8\pm1.5$	18 / 20		





#### XENON10 data (Sorensen 2008)

<u>Model (0 or 1 free parameter):</u> Without e- extraction delay With extraction delay at L/GXe

- \*Can input various electron effects into simulation to get increasingly accurate picture of S2 signal shape in time (NEST aims for full detector sim, for everyone)
- \*We can (crudely) model the absolute yield of UV photons produced in the gas gap of a XeTPC per successfully extracted electron vs. field

# SZ PHISE Shape

\*Dark matter is probably there, and there is a fairly agnostic way to directly look for one candidate, WIMPs, by waiting for nuclear recoils to produce a signal, while electron recoils are the main background

\*Two-phase XeTPCs are a great way to look for WIMPs

- \*NEST is an ever-evolving codebase/software and corresponding collection of semi-empirical models that works really well not only for incorporating old data, but for predicting new results *a priori* and designing new experiments with careful simulation
- \*In near future, switching gears to quantum chemistry, atom by atom, closer to first-principles approach





### \*Questions?